

**Key words.** ISM: supernova remnants — X-rays: individual:  
DEM L241 — X-rays: individual: XMMU J053559.3–673509

# A Detailed Observation of a LMC Supernova Remnant DEM L241 with *XMM-Newton*<sup>★</sup>

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**Abstract.** We report on an *XMM-Newton* observation of the supernova remnant (SNR) DEM L241 in the Large Magellanic Cloud. In the soft band image, the emission shows an elongated structure, like a killifish, with a central compact source. The compact source is point-like, and named as XMMU J053559.3–673509. The source spectrum is well reproduced with a power-law model with a photon index of  $\Gamma = 1.57$  (1.51–1.62) and the intrinsic luminosity is  $2.2 \times 10^{35}$  ergs s<sup>−1</sup> in the 0.5–10.0 keV band, with the assumed distance of 50 kpc. The source has neither significant coherent pulsations in  $2.0 \times 10^{-3}$  Hz–8.0 Hz, nor time variabilities. Its luminosity and spectrum suggest that the source might be a pulsar wind nebula (PWN) in DEM L241. The spectral feature classifies this source into rather bright and hard PWN, which is similar to those in Kes 75 and B0540–693. The elongated diffuse structure can be divided into a “Head” and “Tail”, and both have soft and line-rich spectra. Their spectra are well reproduced by a plane-parallel shock plasma (*vpshock*) model with a temperature of 0.3–0.4 keV and over-abundance in O and Ne and a relative under-abundance in Fe. Such an abundance pattern and the morphology imply that the emission is from the ejecta of the SNR, and that the progenitor of DEM L241 is a very massive star, more than 20 M $\odot$ . This result is also supported by the existence of the central point source and an OB star association, LH 88. The total thermal energy and plasma mass are  $\sim 4 \times 10^{50}$  ergs and  $\sim 200$  M $\odot$ , respectively.

## 1. Introduction

Supernovae (SNe) and supernova remnants (SNRs) shape and enrich the chemical and dynamical structure of the interstellar medium and clouds. X-ray studies give us plenty of information about hot plasma in SNRs with emission lines from highly ionized ions. Moreover, SNRs are believed to be cosmic ray accelerators around their pulsar and pulsar wind nebula (PWN), and/or shock fronts. The Magellanic Clouds (MCs) are the best galaxies for the systematic study of SNRs, thanks to the known distance (50 kpc; Feast, 1999) and the small absorption column. Another subject of interest is supernova explosions in starburst galaxies. The SNe of massive stars scatter light elements into such galaxies, and make an abundance pattern different from those in normal galaxies (e.g., Umeda et al., 2002). The Large Magellanic Cloud (LMC) is the nearest starburst galaxy (Vallenari et al., 1996), and thus we can examine the influence of discrete massive SNe on the galaxy. Now, more than 30 LMC SNRs are cataloged

(Williams et al., 1999) and the number may still increase (e.g., Chu et al., 2004). However, we have only few samples for which type and age are known. The pulsars and their nebulae are a clear evidence of core-collapsed origin and are good indicators of their age. Now, only four LMC SNRs are reported to have a pulsar and/or PWN; B0540–693 (Manchester et al., 1993), N157B (Wang et al., 2001), B0532–710 (Klinger et al., 2002; Williams et al., 2005), and B0453–685 (Gaensler et al., 2003). This is mainly because of a lack of spatial resolution in previous radio and X-ray observations. Thus, searching for new SNR containing pulsars and PWNe in the LMC is critical to carrying out a systematic study of the contributions from SNRs to interstellar medium in the LMC.

DEM L241 (0536–67.6) was identified as an SNR by Mathewson et al. (1985). They found that DEM L241 shows the typical looped filamentary structure of an SNR with a size of  $\sim 2'$  in the optical image. The [S II] to H $\alpha$  ratio is 0.6 (Mathewson et al., 1985), which is typical of that for SNRs. Radio emission in 843 MHz was also reported by Mathewson et al. (1985), which shows shell-like structure with

similar size to the optical shell. *Einstein* detected relatively bright X-rays ( $2.7 \times 10^{35}$  ergs s<sup>-1</sup> in the 0.15–4.5 keV band, Mathewson et al., 1985). It appears to be a blow-out of the dense H II region, N 59B, around the OB star association LH 88 (Chu & Kennicutt, 1988). A study of the echelle spectrum by Chu (1997) shows a relatively small expansion velocity and a large intrinsic velocity width, implying that the SNR may be expanding within the stellar-wind blown cavity, which suggests that the progenitor of DEM L241 is likely to be a massive star. Nishiuchi (2001) observed this SNR with *ASCA*, and found that the spectrum of the remnant requires not only thermal emission but also a power-law component with  $\Gamma \sim 1.6$ . This result indicates that DEM L241 has a pulsar and/or a PWN, or synchrotron X-ray emitting shells like SN 1006 (Koyama et al., 1995; Bamba et al., 2003). However, the lack of spacial resolution of *ASCA* prevented us from concluding the origin of the hard X-ray emission.

In this paper, we report the detailed X-ray analysis of DEM L241 for the first time using *XMM-Newton*, with the help of *ASCA* data. §2 summarizes the details of *XMM-Newton* and *ASCA* observations of DEM L241. The analysis results are written in §3. §4 is devoted to the discussion about the origin of the remnant. We assume the distance to the LMC to be 50 kpc (Feast, 1999) in this paper.

## 2. Observations and Data Reduction

*XMM-Newton* (Jansen et al., 2001) EPIC cameras (Strüder et al., 2001; Turner et al., 2001) observed DEM L241 on December 29 2004 (observation ID = 0205380101). The on-axis point spread function (PSF) is 5'' for MOS CCDs and 6'' for pn, in full width half maximum. The medium filter was used for all three EPIC cameras to block ultraviolet photons (Stephan et al., 1996; Villa et al., 1998). In this observation, both EPIC MOS and pn cameras were operated in the full frame mode providing a time resolution of 2.6 s and 73.4 ms, respectively. The data reductions and analysis were made using Science Analysis System (SAS) software version 6.1.0. We filtered out background flares for rates in >10 keV band more than 0.35 (1.0) counts s<sup>-1</sup> for the MOS (pn) cameras, resulting in 45 (43) ks of good time intervals. In the following analysis, we use the grade pattern 0–12 events for MOS data and 0–4 events for pn data, according to the SAS guide.

For the timing analysis, we used *ASCA* (Tanaka et al., 1994) GIS (Gas Imaging Spectrometer, Ohashi et al., 1996) screened archive data of DEM L241, because GIS has better time resolution than pn. The observation was carried out on October 12–14 1999. Two GISs (GIS2 and GIS3) were operated in the nominal pulse height mode. Only high bit data was used, which provides the best time resolution of 62.5 ms. The total exposure time is 42 ks for both GISs. To increase the statistics, the data of the two detectors, GIS2 and GIS3, were combined in the following study.

## 3. Results

### 3.1. Image Analysis

Figure 1 shows the *XMM-Newton* MOS 1+2 images of DEM L241 in the (a) 0.5–2.0 keV and (b) 2.0–9.0 keV bands. The correction of the exposure time was performed, whereas the subtraction of background photons was not carried out. In the soft band image, we can see a diffuse structure elongated from southeast to northwest with the size of  $\sim 1'.5 \times 3'$ , corresponding to 22 pc $\times$ 44 pc at 50 kpc. The shape is like a killifish, with a double peaked feature on its “Head” and “Tail” (see Figure 1). In addition to the body of the fish, there is a compact source like an “eye” of the fish. On the other hand, only the eye can be seen in the hard band image.

We determined the position of the point-like source with the MOS 1+2 image, using the *edetect* command in SAS, to be (05<sup>h</sup>35<sup>m</sup>59<sup>s</sup>.93, -67<sup>d</sup>35<sup>m</sup>09<sup>s</sup>.72), then named the source as XMMU J053559.3–673509. Although counterparts were searched for with the SIMBAD data base, we found no candidate in any wavelength. Hereafter, we refer to this source as Source 1. The radial profile of Source 1 was made with the 2.0–9.0 keV MOS 1+2 data as shown in Figure 2. The source is as compact as the PSF of *XMM-Newton*. The fitting of the profile with a single King profile model was accepted statistically with reduced  $\chi^2$  of 18.0/17. The best-fit core radius is 3''.1 (2''.4–3''.9) (hereafter, the parentheses indicate single parameter 90% confidence regions), which is same as the PSF of MOS CCDs. Therefore, we concluded that Source 1 is not significantly extended. The upper-limit of the source size is 1.0 pc.

Figure 3 shows the [S II] gray scale image (Mathewson et al., 1985), overlaid on the MOS 1+2 0.5–9.0 keV contour map. The [S II] emission, which traces the shock front, surrounds the Head region with two strong and weak rims on northeastern and southwestern sides. The 843 MHz map by Mathewson et al. (1985) basically agrees with the [S II] map. The northern part of the Head is also bright in the light of [O III] (Mathewson et al., 1985). There is no enhancement around the Tail region in both [S II] and radio maps, although the field of view of the [S II] map is too small to cover the whole Tail region.

### 3.2. Spectral Analysis

#### 3.2.1. Source 1

The spectrum of Source 1 was accumulated from a 15'' radius circle around the source for each detector. The background annular region was selected with 20'' inner radius and 30'' outer radius. Figure 4(a) shows the background subtracted spectrum of Source 1, which is very hard and has no line-like structure. We fitted the spectra of the three detectors simultaneously with a power-law function plus absorptions. The galactic and LMC absorptions ( $N_{\text{H}}^{\text{gal}}$  and  $N_{\text{H}}^{\text{LMC}}$ , respectively) were calculated separately. We fixed  $N_{\text{H}}^{\text{gal}}$  to be  $5.56 \times 10^{20}$  cm<sup>-2</sup>, which is estimated by Dickey & Lockman (1990), whereas  $N_{\text{H}}^{\text{LMC}}$  was treated as a free parameter. Each absorption column was subsequently calculated using the cross sections by Morrison & McCammon (1983) with the solar abundances

(Anders & Grevesse, 1989) for  $N_{\text{H}}^{\text{gal}}$ , and the cross sections by Balucinska-Church & McCammon (1992) with the average LMC abundances (0.3; Russell & Dopita, 1992; Hughes et al., 1998) for  $N_{\text{H}}^{\text{LMC}}$ . The fitting was acceptable with the reduced  $\chi^2$  of 365.1/376. The best-fit model and parameters are shown in Fig. 4(a) and Table 1.

### 3.2.2. Diffuse emission

The diffuse emission in DEM L241 shows an elongated structure with a “Head” and “Tail”. We divided the emission into the Head and Tail regions as shown in Figure 1(a). The source region for the spectral analysis of the Head was made from a  $70''$  radius circle region excluding a  $30''$  radius circle around Source 1. On the other hand, the region for the Tail was taken as an elliptical shape with  $50'' \times 90''$  radii. The background photons were accumulated from a source free region near the SNR. These regions are shown in Figure 1(a), with thick (for the sources) and thin (for the background) lines.

Figure 4(b) and (c) show the background-subtracted spectra of the Head and Tail regions. The spectra are basically soft and have line-like structure, implying that there is at least a thermal emission component. The center energies of these lines agree with He-like O, He-like Ne, and He-like Mg. The Head spectrum has an additional hard tail, which is probably contamination from the bright Source 1, since the fractional encircled energy within  $30''$  is only  $\sim 90\%$ .

We used the spectra of the three CCDs simultaneously and fitted them with a plane-parallel shock plasma (*vpshock*) model (Borkowski et al., 2001), packaged in *xspec* 11.3.1., with the galactic and LMC absorptions. The absorption calculation has been done in the same way to Source 1 case (see §3.2.1). The abundances of the plasma were fixed to be 0.3 (Russell & Dopita, 1992; Hughes et al., 1998). For the spectrum of the Head region, we added an additional power-law component which represents contamination from Source 1. The value of  $\Gamma$  was fixed to be 1.57, the best-fit value for Source 1 (Table 1). The fittings were rejected with the  $\chi^2/\text{d.o.f.}$  of 871.8/381 (Head) and 607.1/300 (Tail), with sinusoidal residuals around O and Ne lines. Therefore, we allowed O, Ne, and Fe abundances to vary freely following the previous result by Hughes et al. (1998). The fits were greatly improved, and accepted statistically for both regions with reduced  $\chi^2$  of 505.3/377 (Head) and 333.9/296 (Tail). The best-fit models and parameters are shown in Figure 4 (b) and (c), and Table 2, respectively. We tried to but gave up being free the abundance of Si, because there is a strong background emission line (Katayama et al., 2004).

The flux of power-law component in the Head region is  $6.4 \pm 1.3\%$ , which is consistent with the expected contamination of Source 1. Thus, we concluded that this hard excess is due to the contamination of Source 1 only, and do not mention this component hereafter.

### 3.3. Timing Analysis

Coherent pulsations were searched for in Source 1. We extracted photons from the 2.0–8.0 keV band (Source 1 is bright in this band as seen in Figure 4(a)) and from a  $15''$  radius circle around Source 1. The correction of the photon arrival times was done to those at the barycenter. A Fast Fourier Transform (FFT) algorithm was applied to the pulsation search. We used all three instruments for the pulsation search from  $2 \times 10^{-3}$  Hz to 0.19 Hz, and only the pn CCD, which has better time resolution than the MOS CCDs, for the 0.19–6.8 Hz search. The resulting power density spectra (Figure 5a, b) show no significant peak from  $2 \times 10^{-3}$  Hz to 6.8 Hz. Although we also analyzed timing data accumulated from a  $5''$  radius circle and/or photons in other wider and narrower bands, the results did not change significantly.

We also conducted the timing analysis with ASCA GIS data, which has better time resolution than MOS and pn. The source photons were accumulated from a  $3'$  radius circle around Source 1, within the 2.0–8.0 keV band. After correcting the photon arrival times to those at the barycenter, we applied a FFT algorithm to search the coherent pulsation. No significant pulsation was found up to 8.0 Hz again, as can be seen in Figure 5(c).

Figure 6 shows the light curve of Source 1 in the 2.0–8.0 keV band, which is the combined result of the three *XMM-Newton* detectors. We have done the Kolmogorov-Smirnov test for this light curve and found that Source 1 has no significant time variability during the observation with the probability of constancy of 0.72.

## 4. Discussion

### 4.1. The absorption

The best-fit  $N_{\text{H}}^{\text{LMC}}$  in our analysis is almost consistent with or slightly larger than the LMC absorption column estimated from the 21 cm line survey ( $8.4 \times 10^{20} \text{ cm}^{-2}$ ; Rohlfs et al., 1984), implying that this source is in the LMC. The small excess might be due to the H I excess reported in the H I map by Staveland-Smith et al. (1998). However, we can conclude nothing due to the lack of high sensitivity information around the SNR. More H I observations with better spatial resolution are required to discuss this issue quantitatively.

### 4.2. Source 1

The central position of Source 1, the hard spectrum, and persistent luminosity indicate that Source 1 might be a pulsar and/or a PWN of DEM L241. The lack of detection of coherent pulsation is not surprising, since the periods of ordinary rotation-powered pulsars detected with X-rays are  $\leq 0.1$  sec. The hard spectrum also resembles the emission from background active galactic nuclei (AGNs), although the source shows no time variability. We estimate the probability of chance coincidence that there is a background AGN accidentally in the SNR region. With the  $\log N - \log S$  relation of AGNs derived by Hasinger et al. (1998), we found that the expected number of AGNs is only  $1.1 \times 10^{-3}$  in the  $1.5' \times 3'$  region. Therefore, we

concluded that Source 1 is an LMC member. The similarity of the absorption column for Source 1 to that of the Head region also support that. In the case that Source 1 is a pulsar and/or a PWN, this adds a new member to the sample of LMC PWNe (Gaensler et al., 2003). The hard X-ray flux detected by ASCA (Nishiuchi, 2001) is consistent with that of Source 1. Thus we consider that is not nonthermal emission from energetic electrons accelerated on the shell, but the central point source. The distance from Source 1 to the center of LH 88 is about  $14''$ , and the diaphragm diameter of LH 88 is  $100''$  (Bica et al., 1996), so the progenitor might be a member of the cluster.

The intrinsic luminosity of Source 1 is  $2.2 \times 10^{35}$  ergs  $s^{-1}$  in the 0.5–10.0 keV band under the assumption that the distance to Source 1 is 50 kpc, and the photon index is 1.57 (1.51–1.62). The luminosity and hardness ranges of pulsars are wide,  $\log L_x \sim 31\text{--}36$ ,  $\Gamma \sim 0.6\text{--}2.0$ , respectively (e.g., Gotthelf & Olbert, 2001; Gaensler et al., 2003). Recently, “compact central sources” have been discovered in several SNRs (Cas A, Chakrabarty et al. 2001; Vela Jr., Kargaltsev et al. 2002; Kes 79, Seward et al. 2003), which have rather small luminosities ( $\log L_x \sim 32\text{--}34$ ) and soft spectra ( $\Gamma = 3\text{--}4$ ). In the case that Source 1 is a pulsar or a compact central source, Source 1 is one of the brightest samples, and as bright as the Crab pulsar (Pravdo et al., 1997). Such bright pulsars certainly have PWNe, which are about 10 times brighter than the pulsars themselves and also have hard spectra ( $\log L_x = 33\text{--}37$  and  $\Gamma = 1.3\text{--}2.3$ ; Gotthelf & Olbert, 2001). Thus, we conclude that the emission is from a PWN of DEM L241. Its photon index is rather small for ordinary PWNe, and similar to those of the PWN in Kes 75 (Helfand et al., 2003) and B0540–693 (Hirayama et al., 2002). Helfand et al. (2003) notes that these pulsar systems change their energy into X-ray radiation very efficiently. Source 1 may be a new sample of such energetic PWNe. Source 1 is moderately bright for PWNe, which might imply that the system is middle-aged. Pulsars in such PWNe have periods of less than a few hundred msec and luminosities of about 10% of their nebulae. Therefore it is natural that we could not detect the coherent pulsation even in the case that Source 1 has a pulsar.

In order to confirm the origin of Source 1, it is essential to resolve the pulsar spatially and to detect coherent pulsations. Therefore, further observations are encouraged with better time and spatial resolution, and excellent statistics, in the X-ray and radio bands.

#### 4.3. Diffuse emission

The temperature and abundance pattern of the hot plasma in the Head and Tail regions are roughly statistically the same as each other, (Table 2), implying that these plasmas have the same origin. The overabundant O and Ne relative to the average LMC indicate that the plasma emission is from the ejecta of the explosion, and that relative to Fe implies that the progenitor of DEM L241 may be a core-collapsed SN (Tsujiimoto et al., 1995). The Fe abundance is slightly smaller than the average value in the LMC (0.22 Hughes et al., 1998), it might be the local fluctuation of metal. The less abun-

dant O relative to Ne might imply that the progenitor is very massive,  $\geq 20M_{\odot}$  (Umeda et al., 2002). Umeda et al. (2002) also insists that fast convective mixing in the progenitor star (Spruit, 1992) makes more Ne and less O. Therefore, we concluded that the progenitor of DEM L241 was very massive star. This conclusion is also supported by the presence of the energetic pulsar candidate and the OB star association, LH 88 (Chu & Kennicutt, 1988). The abundance pattern of DEM L241 resembles in that of a starburst galaxy M82 (Tsuru et al., 1997). This fact might imply that DEM 241 have been produced in a period of starburst activity of the OB star association. Umeda et al. (2000) and Nakamura et al. (2001) have shown that a large Si/O ratio may be signature of energetic SN explosion ( $\geq 10^{51}$  ergs). Thus, we need to determine Si abundance with low background observations.

The thermal emission in the Head region has a center-filled morphology even considering the contribution from Source 1. Note that the radius of the extent of the thermal emission is  $\sim 70''$  while that of Source 1 corresponds to the *XMM-Newton* PSF,  $\sim 5''$ . The structure, the radio shell (Mathewson et al., 1985), and the large size of the SNR (see §3.1), classify the SNR as a mixed-morphology (MM) SNRs (Rho & Petre, 1998). However, the over-abundant light elements suggest that the emission is not from interstellar medium but ejecta as already mentioned, which is not common for ordinary MM SNRs. Then, we concluded that this SNR is an ejecta dominant SNR without X-ray emitting shells. The absence of a limb-brightened X-ray shell is expected, since this SNR is in a hot and tenuous bubble (Chu, 1997) and it is hard to form the strong shock. The elongated shape of the Tail region, together with the lack of enhancement of [S II] and radio continuum emission around Tail, might indicate that the plasma was ejected an-isotropically into a low-density region. It is natural to consider that the plasma in the Tail region is a blown-out on the basis of the [S II] morphology. Chu (1997) suggested that the SNR could be interacting with the inner walls of a bubble to produce the SNR signature. In such a case, the interacting region is heated and emits thermal X-rays, and the emission certainly traces the shock region, like the [S II] emission. If the thermal emission is, on the other hand, center-filled, then we conclude that the thermal X-rays are not from the interacting region but from the ejecta.

In order to estimate physical parameters of the plasma, we assumed that the plasma in the Head region is distributed in a sphere with the radius of  $70''$ , excluding a spherical region with the radius of  $30''$  around Source 1 (total volume  $V_{\text{Head}} = 5.6 \times 10^{59}$  cm $^3$ ), and that the Tail region has an ellipsoid distribution with the radii of  $50'' \times 50'' \times 90''$  ( $V_{\text{Tail}} = 4.0 \times 10^{59}$  cm $^3$ ), and calculated the mean electron density ( $n_e$ ), the age of the plasma ( $t_p$ ), the thermal energy ( $E \approx 3n_e V kT$ ), and the total mass of the plasma ( $M = n_e m_p V$ , where  $m_p$  is the proton mass). The results are summarized in Table 3. The  $t_p$  indicates that the SNR is rather aged, which is consistent with the large size of the SNR.

## 5. Summary

We have conducted a detailed analysis of DEM L241 with *XMM-Newton* for the first time. A summary of our results is as follows:

1. The X-ray emission shows an elongated morphology with a Head and Tail like a killifish. Its size is about  $22 \text{ pc} \times 44 \text{ pc}$  at 50 kpc. The [S II] and radio emission surround the Head, whereas there is no enhancement around the Tail.
2. We resolved a central point source, XMMU J053559.3–673509, in DEM L241 for the first time. The source has no counterpart in any wavelength, and neither coherent pulsation nor time variability is found. The spectrum is well reproduced by a power-law model with parameters of  $\Gamma = 1.57$  and the X-ray luminosity of  $2.2 \times 10^{35}$  ergs in the 0.5–10.0 keV band, implying that this source is a bright hard PWN in DEM L241.
3. The Head and Tail have soft and line-rich spectra with an over-abundance of O and Ne. The abundance pattern indicates a core-collapse origin from a massive progenitor heavier than  $20 M_{\odot}$  for DEM L241.

**Acknowledgements.** This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. Our particular thanks are due to R. M. Williams, M. Nishiuchi, J.S. Hiraga, A. Kubota, I. Takahashi, K. Makishima, P. Ranalli, M. Nakajima, and S. Park, for their fruitful discussions. M. U. and H. N. are supported by JSPS Research Fellowship for Young Scientists. This work is supported in part by the Grant-in-Aid for Young Scientists (B) of the Ministry of Education, Culture, Sports, Science and Technology (No. 17740183).

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**Table 1.** Best-fit parameters for Source 1<sup>a</sup>

Parameters	
$\Gamma$ .....	1.57 (1.51–1.62)
$N_H^{\text{LMC}} [10^{21} \text{cm}^{-2}]^c$ .....	3.3 (2.7–4.0)
Flux [ $10^{-13} \text{ergs cm}^{-2} \text{s}^{-1}$ ] <sup>c</sup> ....	6.4 (6.0–6.8)

<sup>a</sup>: Parentheses indicate single parameter 90% confidence regions.

<sup>b</sup>: Calculated using the cross sections by Morrison & McCammon (1983) with the average LMC abundances (Russell & Dopita, 1992).

<sup>c</sup>: In 0.5–10.0 keV band.

**Table 3.** Physical parameters of the plasma in Head and Tail<sup>a</sup>

Parameters	Head	Tail
$n_e [\text{cm}^{-3}]$ .....	0.28 (0.26–0.29)	0.27 (0.23–0.30)
$t_p [10^6 \text{yrs}]$ .....	(>1.8)	1.1 (>0.2)
$E [10^{50} \text{ergs}]^b$ .....	2.9 (2.2–3.7)	0.6 (0.5–0.7)
$M [M_\odot]$ .....	132 (122–137)	89 (76–99)

<sup>a</sup>: Parentheses indicate single parameter 90% confidence regions.

<sup>b</sup>: Thermal energy  $E$  is estimated to be  $3n_e V kT$ .

**Table 2.** Best-fit parameters for the diffuse emission<sup>a</sup>

Parameters	Head	Tail
<i>vpshock</i>		
$kT_e [\text{keV}]$ .....	0.39 (0.38–0.40)	0.35 (0.34–0.39)
$n_e t_p [10^{12} \text{cm}^{-3} \text{s}]$ .....	(>14.6)	9.0 (>2.0)
$[\text{O}/\text{H}]^b$ .....	0.36 (0.30–0.47)	0.47 (0.36–0.64)
$[\text{Ne}/\text{H}]^b$ .....	0.82 (0.74–0.96)	0.79 (0.63–1.0)
$[\text{Fe}/\text{H}]^b$ .....	0.09 (0.07–0.11)	0.11 (0.09–0.15)
$E.M. [10^{58} \text{cm}^{-3}]^c$ .....	4.4 (3.8–4.8)	2.8 (2.1–3.5)
$N_H^{\text{LMC}} [10^{21} \text{cm}^{-2}]^d$ .....	5.6 (4.7–6.4)	1.8 (1.0–3.1)
Flux [ $10^{-13} \text{ergs cm}^{-2} \text{s}^{-1}$ ] <sup>e</sup> .	3.2	3.6
<i>Power-law</i> <sup>f</sup>		
Flux [ $10^{-14} \text{ergs cm}^{-2} \text{s}^{-1}$ ] <sup>e</sup> .	4.6 (3.6–5.6)	—

<sup>a</sup>: Parentheses indicate single parameter 90% confidence regions.

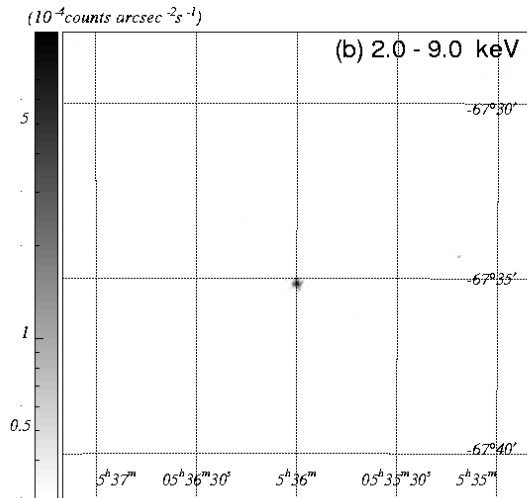
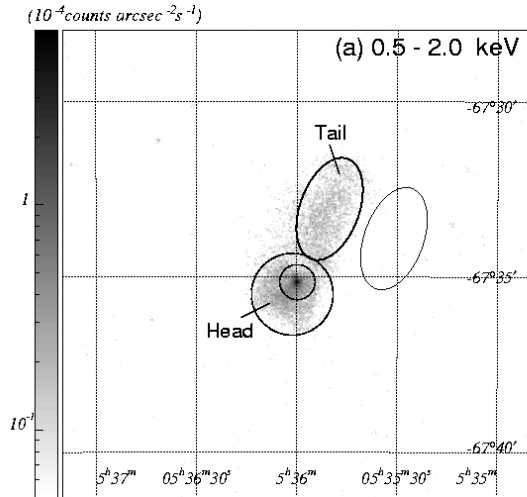
<sup>b</sup>: Abundance ratio relative to the solar value (Anders & Grevesse, 1989).

<sup>c</sup>:  $E.M. = n_e^2 V$ , where  $n_e$  and  $V$  are the electron density and the volume, respectively.

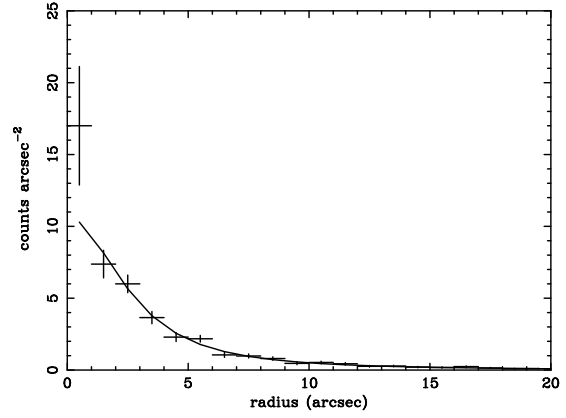
<sup>d</sup>: Calculated using the photometric absorption cross-sections by Balucinska-Church & McCammon (1992) with the LMC abundances (Russell & Dopita, 1992).

<sup>e</sup>: In the 0.5–10.0 keV band.

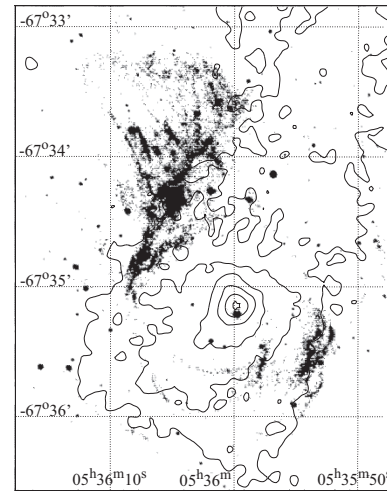
<sup>f</sup>:  $\Gamma$  is fixed to the best-fit value for Source 1 (1.57).



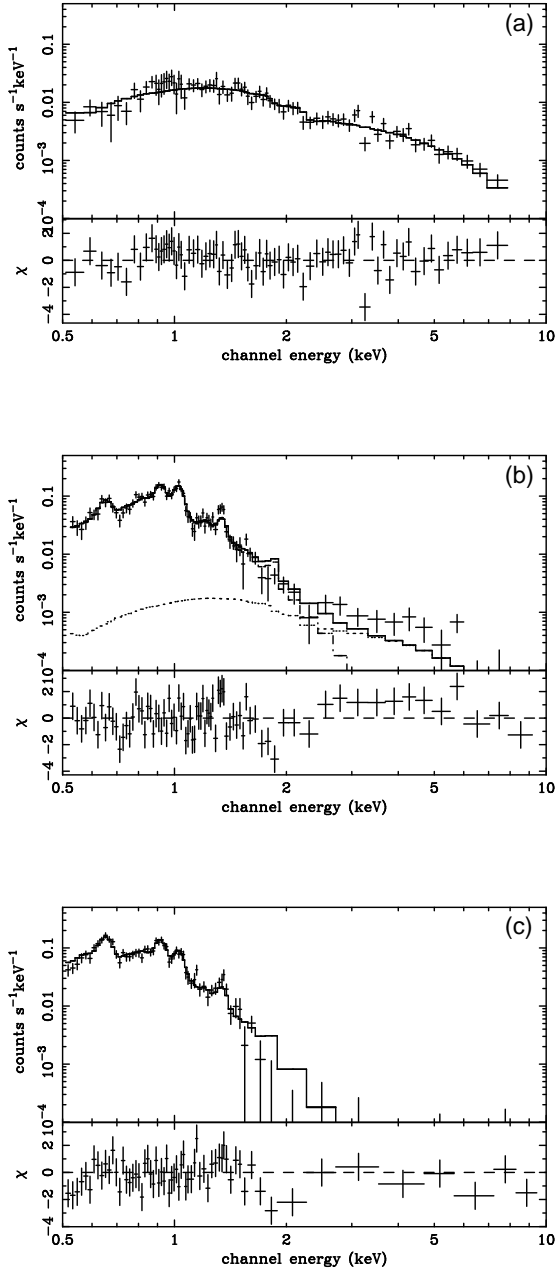
**Fig. 1.** MOS 1+2 images in the (a) 0.5–2.0 keV and (b) 2.0–9.0 keV bands, with J2000 coordinates. The scales are logarithmic, with units of  $\times 10^{-4}$  counts  $s^{-1} \text{arcsec}^{-2} \text{cm}^{-2}$ , as shown in the left bar for each image. Each picture is binned with 20 pixels each and smoothed with  $1''$  scale. The source and background regions used for the spectral analyses are shown by the solid thick and thin lines in the panel (a).



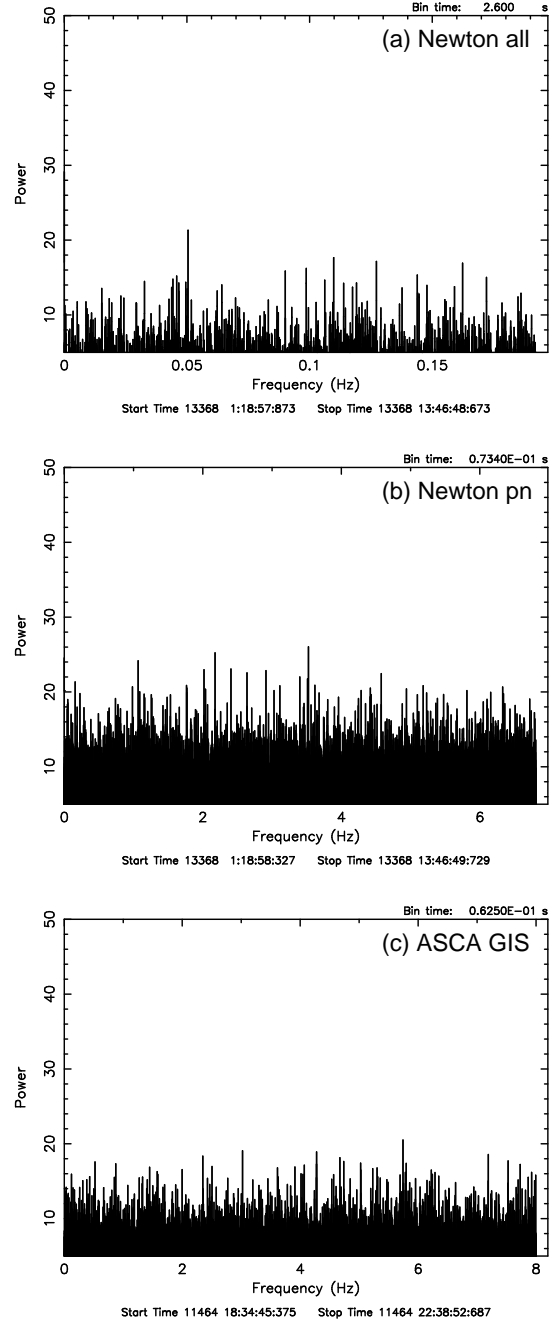
**Fig. 2.** MOS 1+2 radial profile of Source 1 in the 2.0–9.0 keV band. The solid curve and crosses represent the best-fit King profile model and data, respectively.



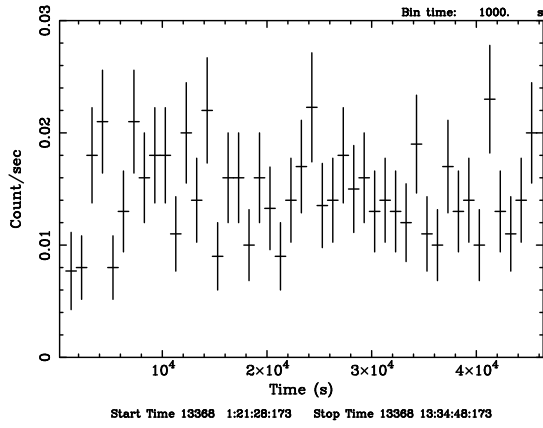
**Fig. 3.** The [S II] gray-scale map (Mathewson et al., 1985), overlaid on the MOS 1+2 0.5–9.0 keV contour map with J2000 coordinates. The contour is in the logarithmic scale.



**Fig. 4.** Spectra of Source 1 (a), Head (b), and Tail (c) regions. Although we fitted the spectra of MOS 1+2 and pn CCDs simultaneously, only the MOS 1 data and results are shown for brevity. The best-fit models are shown with solid (*vpshock*) and dotted (power-law) lines. Lower panels in figures represent data residuals from the best-fit models.



**Fig. 5.** The power density spectra with the MOS1+2 and pn (a), pn (b), and ASCA GIS2+3 (c). Data points are plotted only when whose power is larger than 5.



**Fig. 6.** X-ray light curves of Source 1 with MOS1+2 and pn CCDs in the 2.0–8.0 keV band (bin time = 1ks).